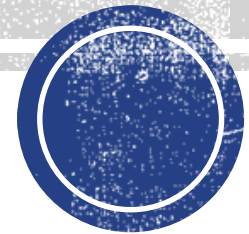
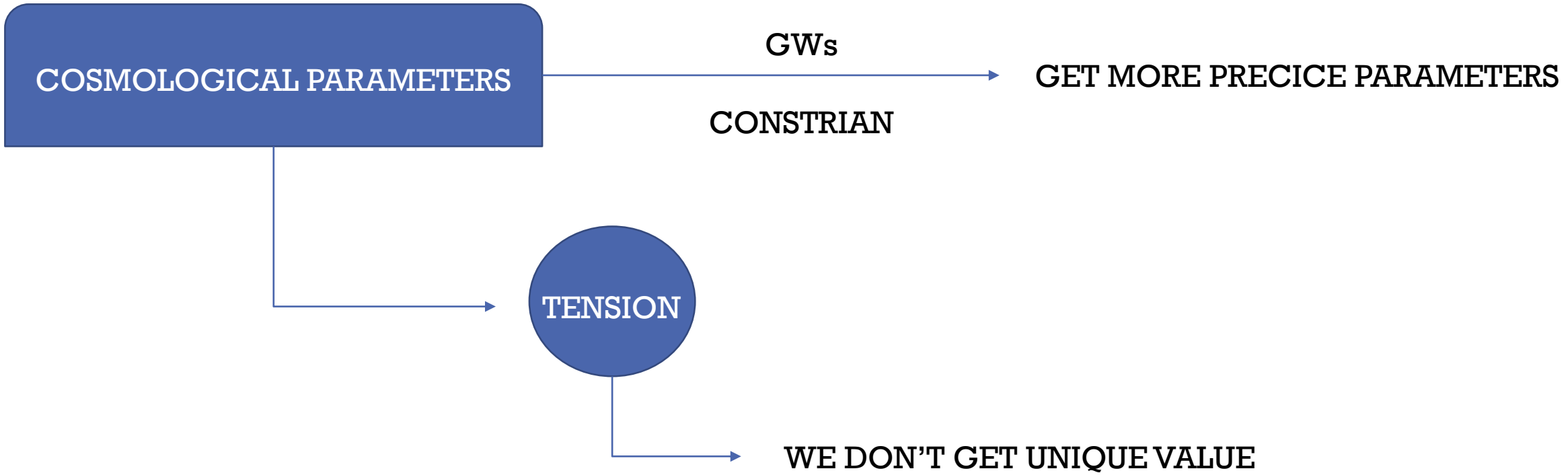
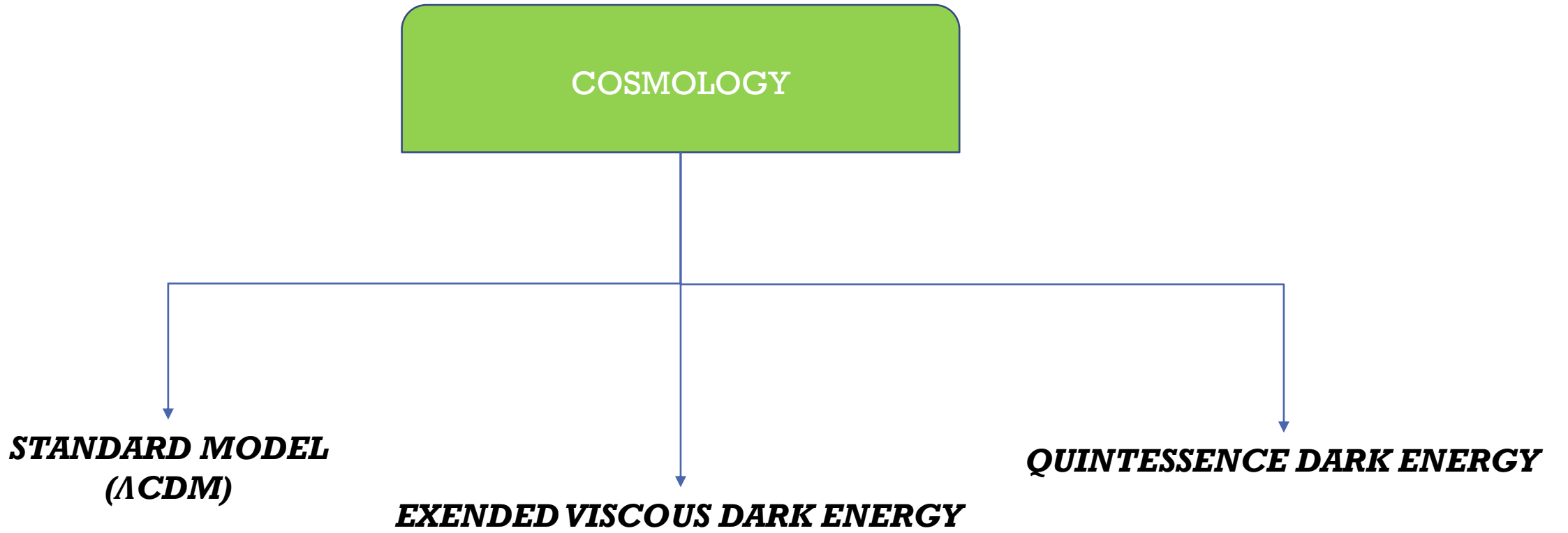


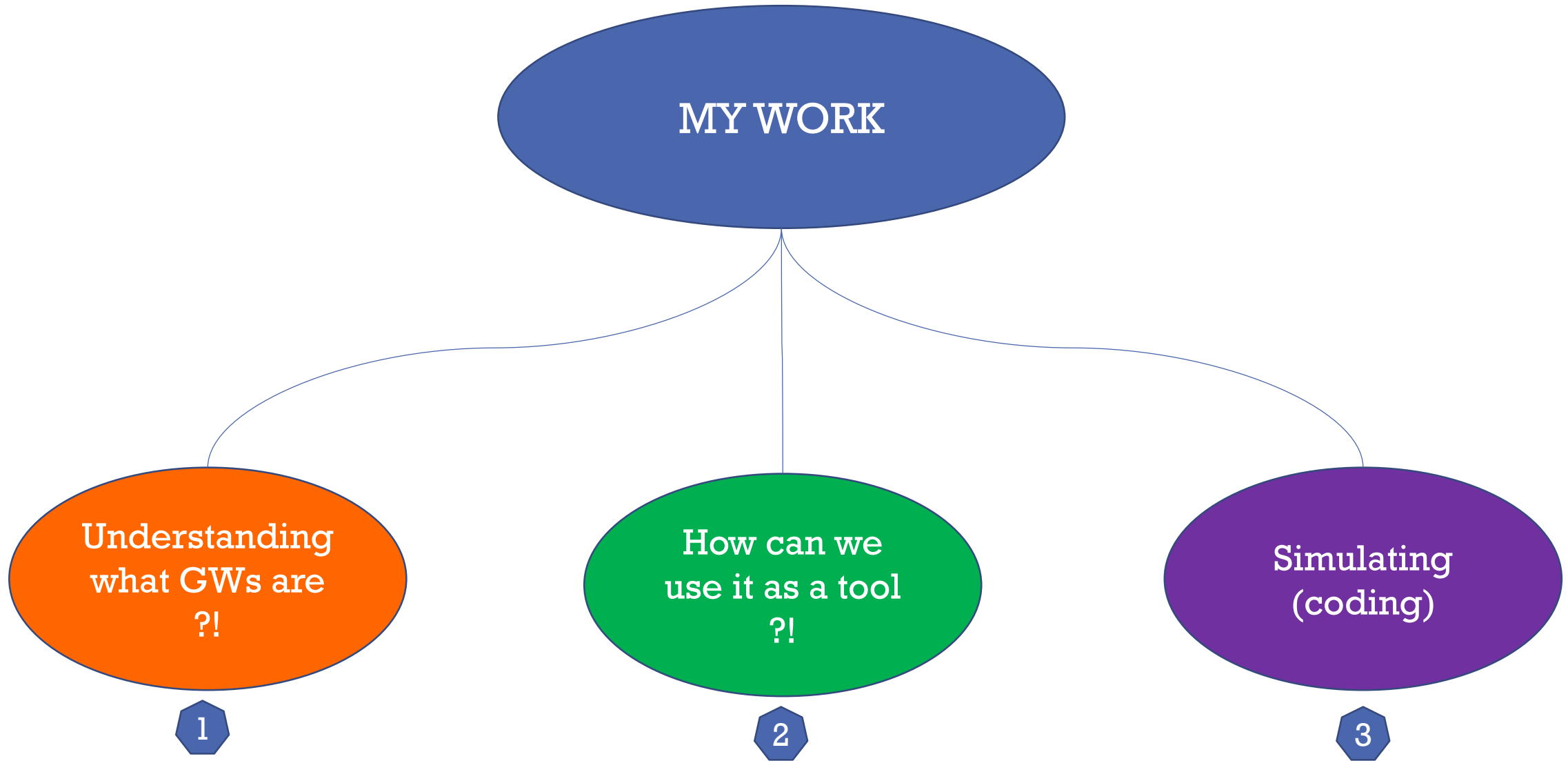
امواج گرانشی به منظور مقیدسازی مدل‌های انرژی تاریک جایگزین  
**GW to constrain alternative dark energy models**



دانشجو : علی غیور  
استاد راهنما: دکتر سیدمحمدصادق موحد

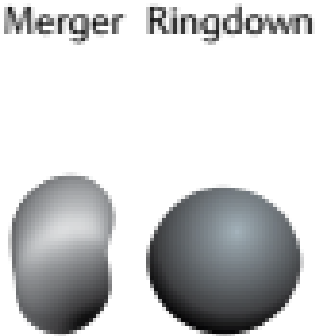
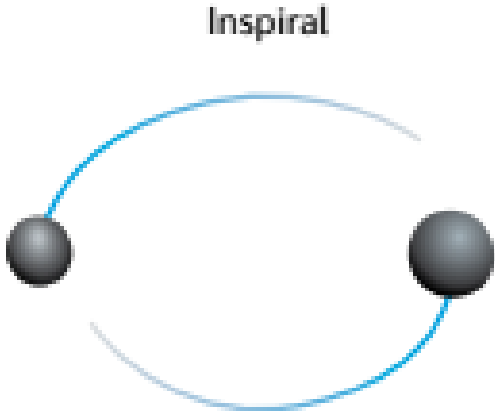
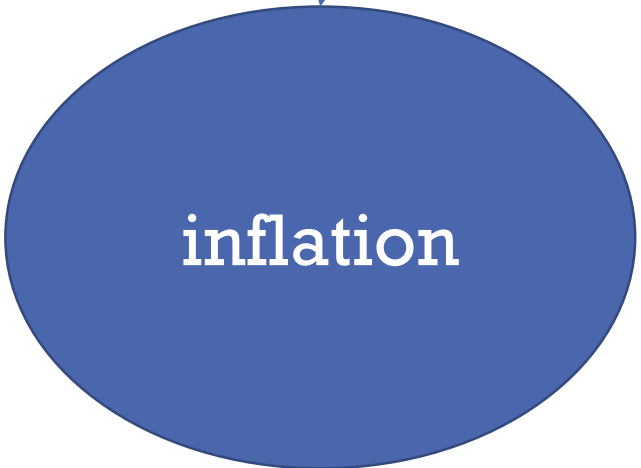






1

# What is GW ?

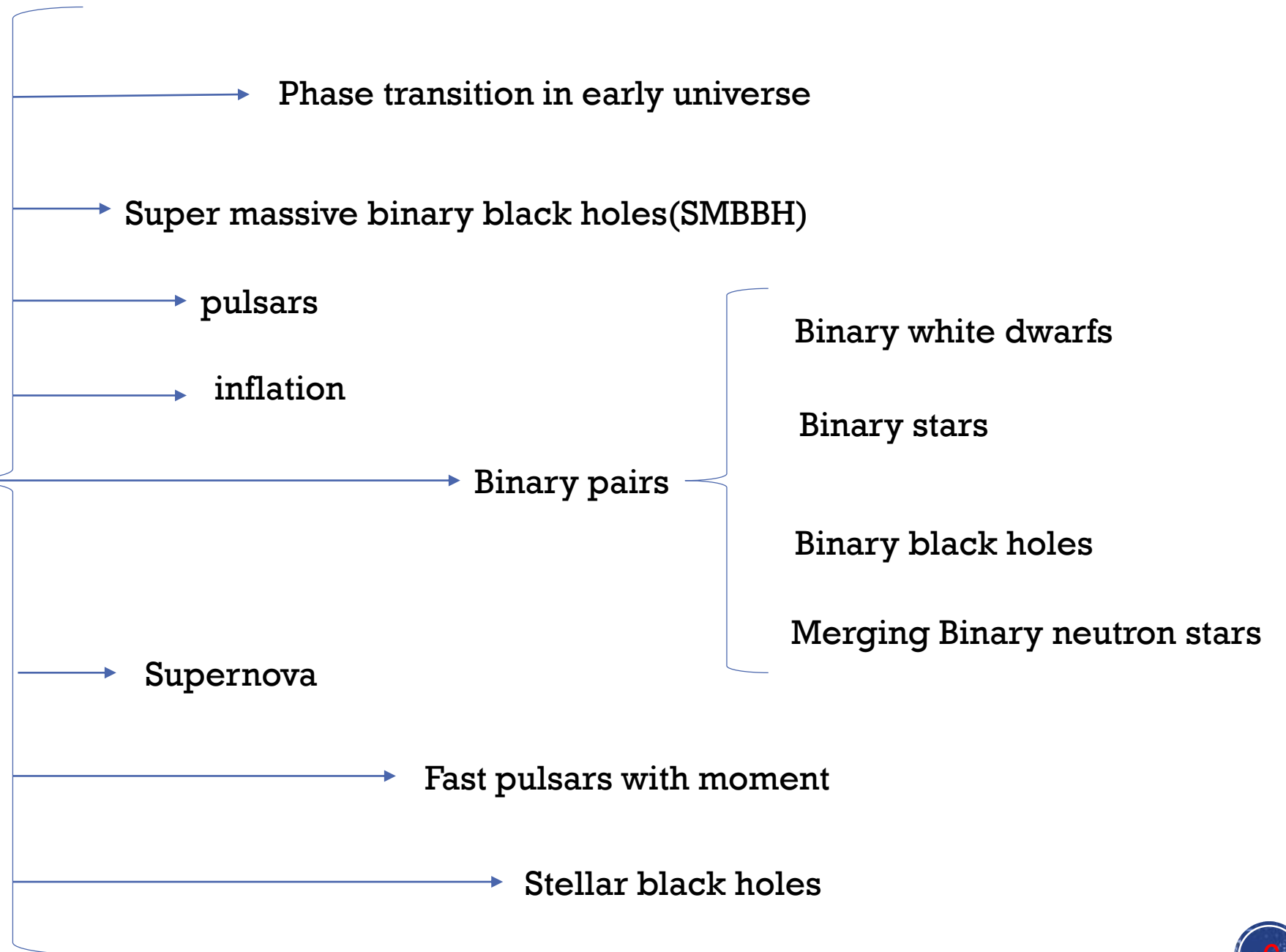


# Why GWs are important for us ?

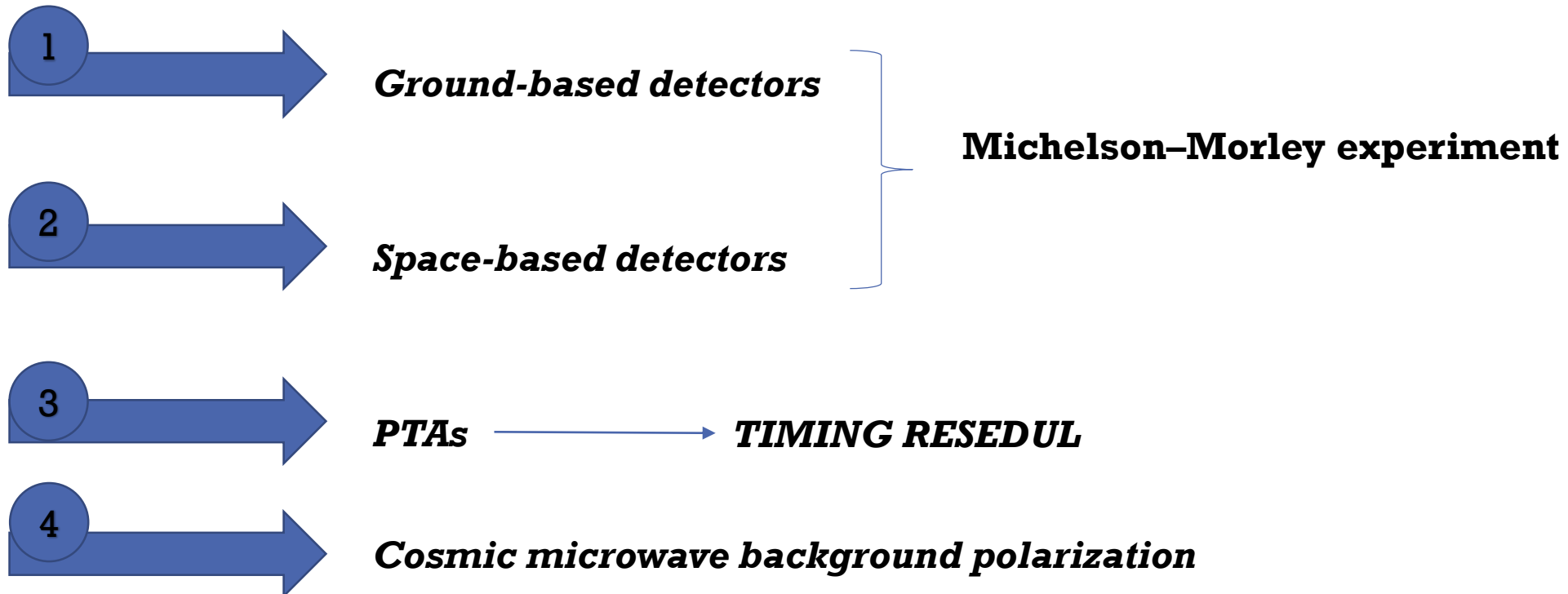
We can understand distance and mass of their sources

We can get more information about EARLY universe

# SOURCES



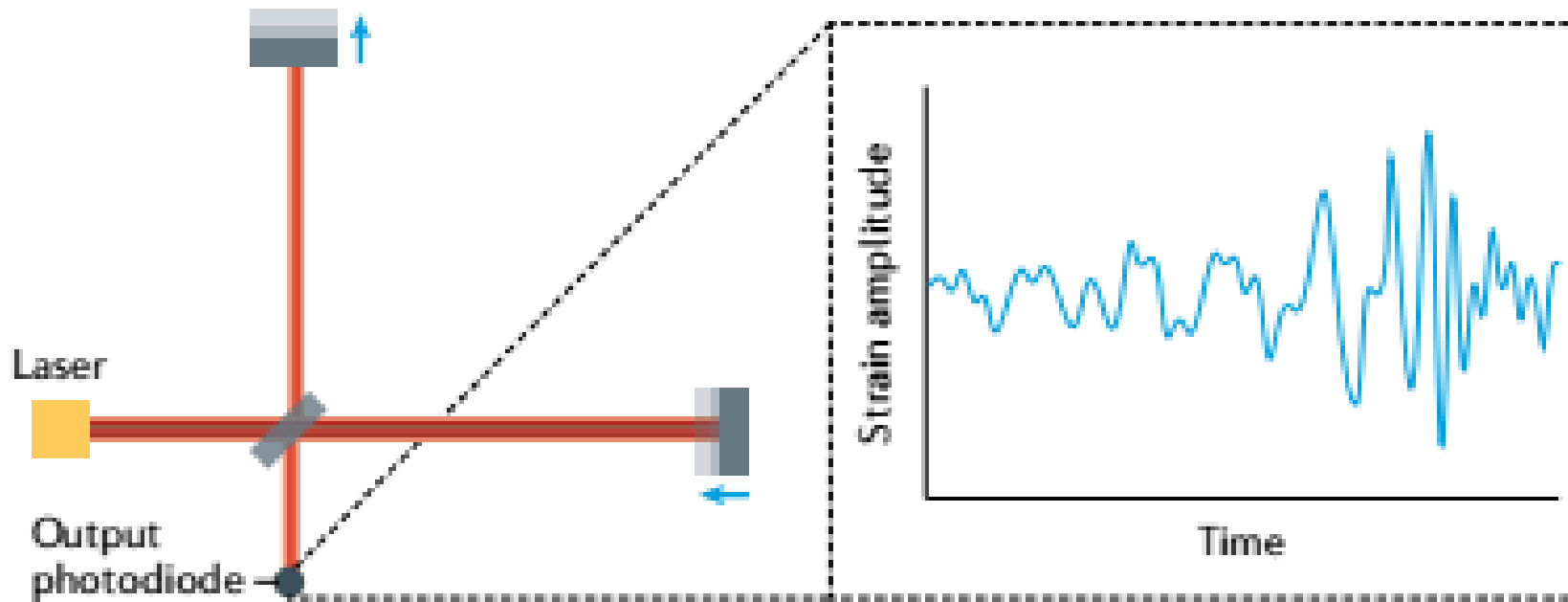
# DETECTORS !





# HOW DO WE DETECT GWs ?

## Michelson–Morley experiment

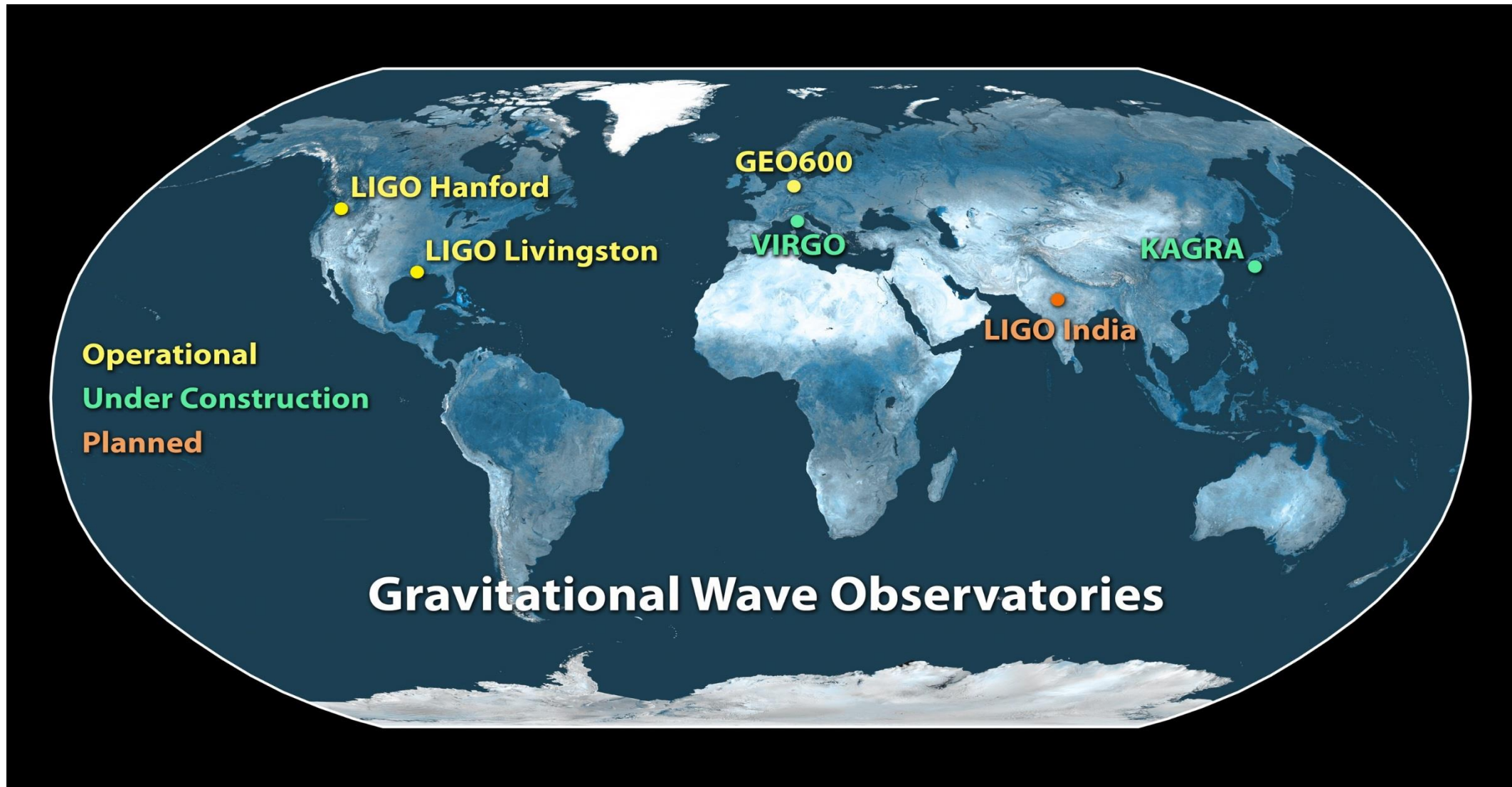


# *Ground-based detectors*

LIGO-VIRGO-KAGRA ...



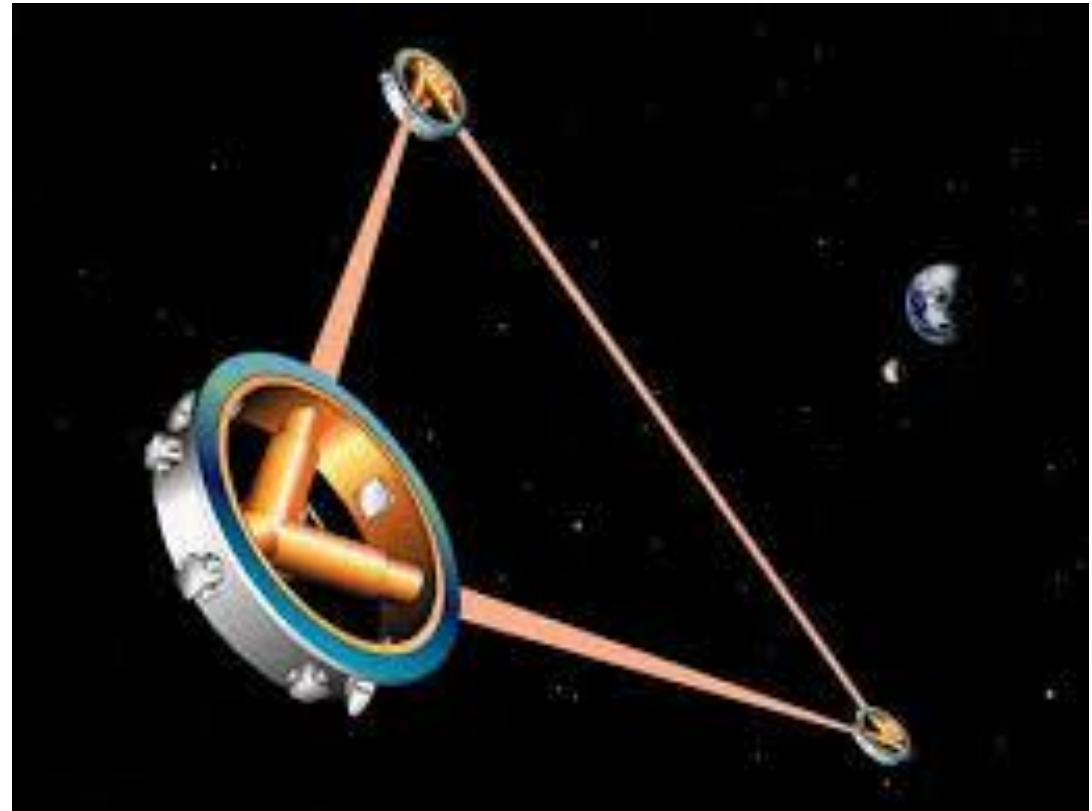
# *Ground-based detectors*



# *Space based detectors*

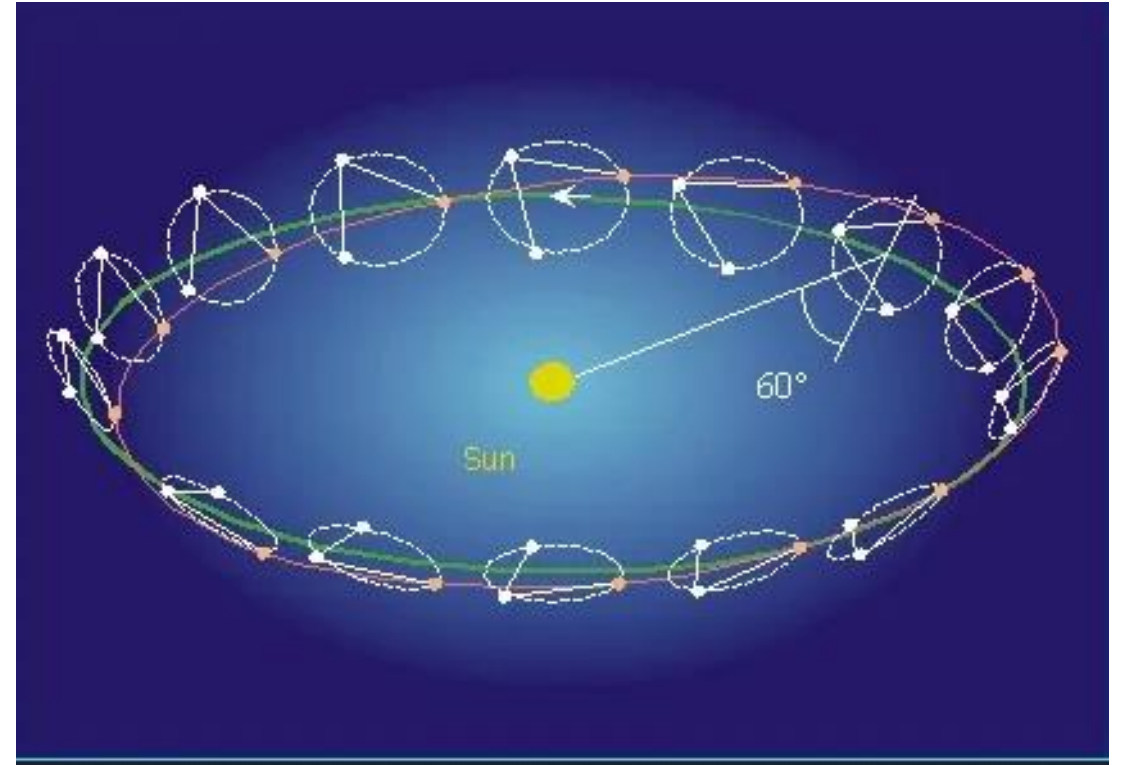
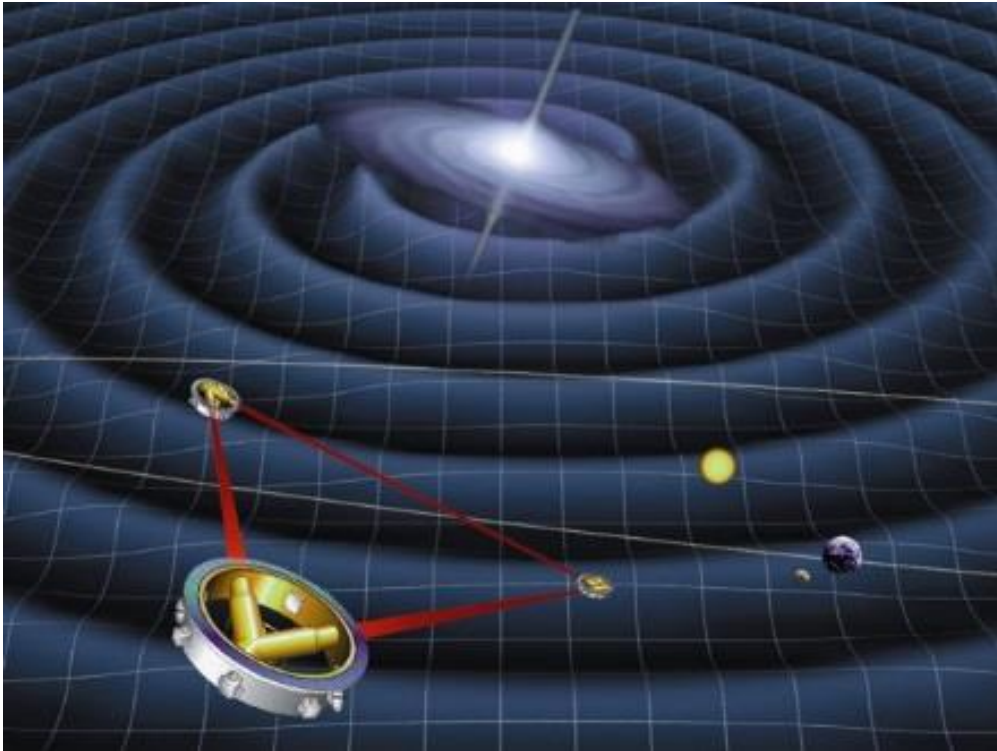
*Laser Interferometer Space Antenna (LISA)*

5Gm

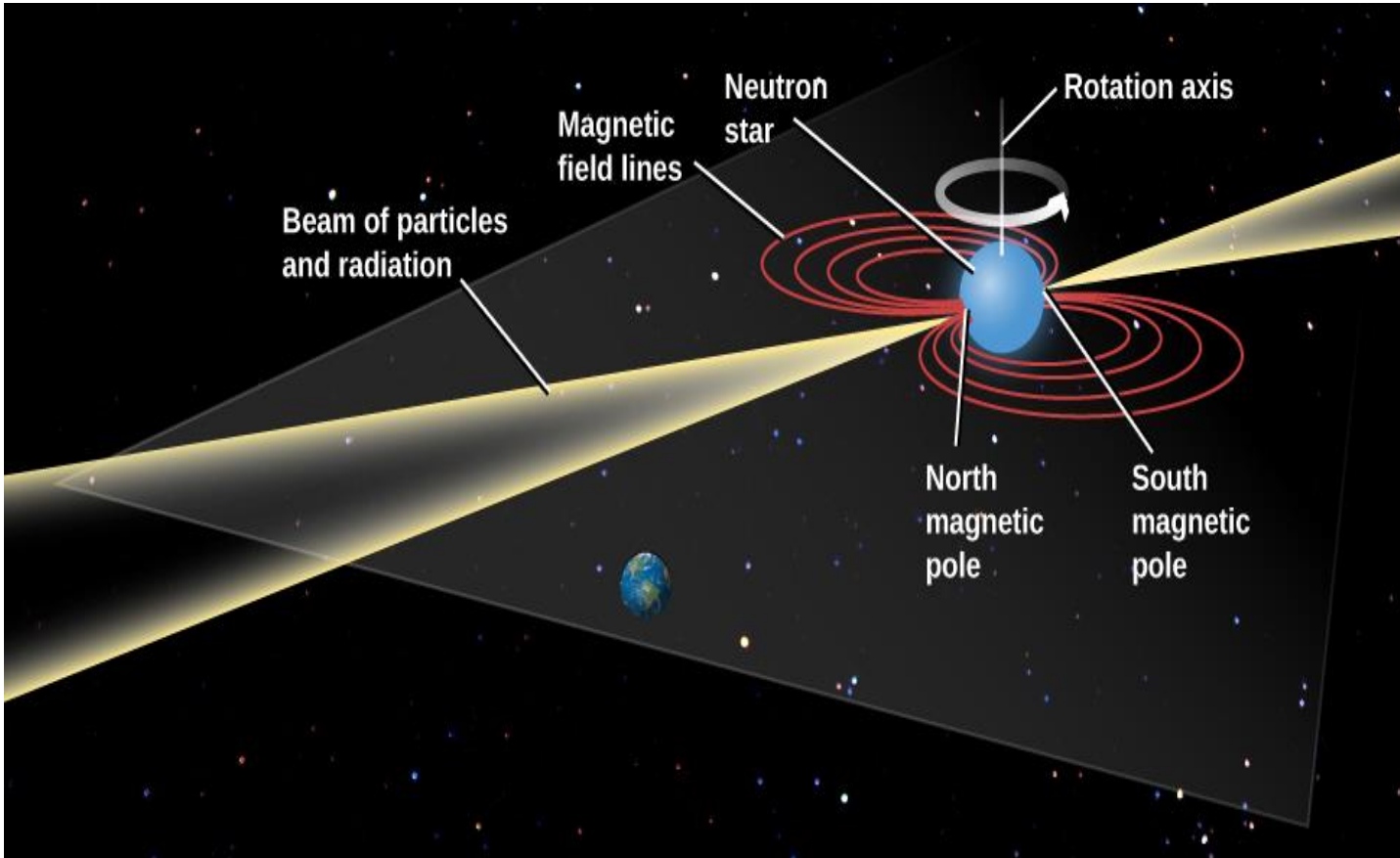


*Deci-hertz Interferometer Gravitational wave Observatory (DECIGO)*

# *Space based detectors*



# *Pulsar timing arrays (PTA)*



Pulsars are rotating neutron stars that act like cosmic lighthouses, appearing as periodic pulsating radio sources. Because millisecond pulsars, pulsars with periods between roughly 1.4 and 30 ms, possess rotational stabilities comparable with the best atomic clocks, they are ideal timing sources.

variations arising from GW perturbations can be measured. Distortions in the spacetime around Earth or the pulsars will produce systematics in timing residuals



**One example**

# Detectors of PTAs



Today, there are three major PTAs: the Parkes PTA<sup>72</sup> in Australia, the European PTA Consortium<sup>65</sup> and the NANOGrav<sup>73</sup> consortium in North America.



# ***Cosmic microwave background polarization***

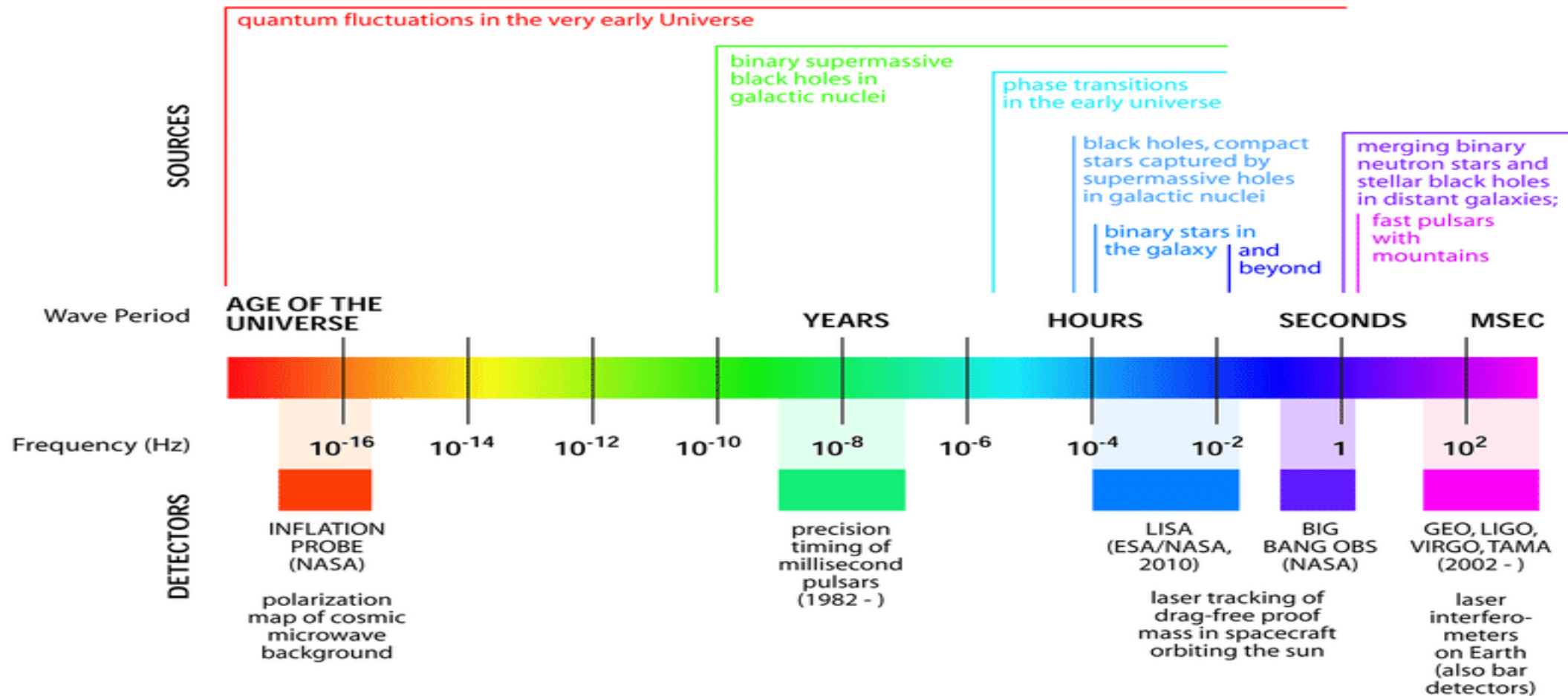
***Laser Interferometer Space Antenna (LISA)***

***Pulsing timing arrays (PTA)***

**Degree Angular Scale Interferometer (DASI)**

# GWs SPECTRUM

## THE GRAVITATIONAL WAVE SPECTRUM



2

What is usage of GWs ?



We can use it to constrain cosmological parameters

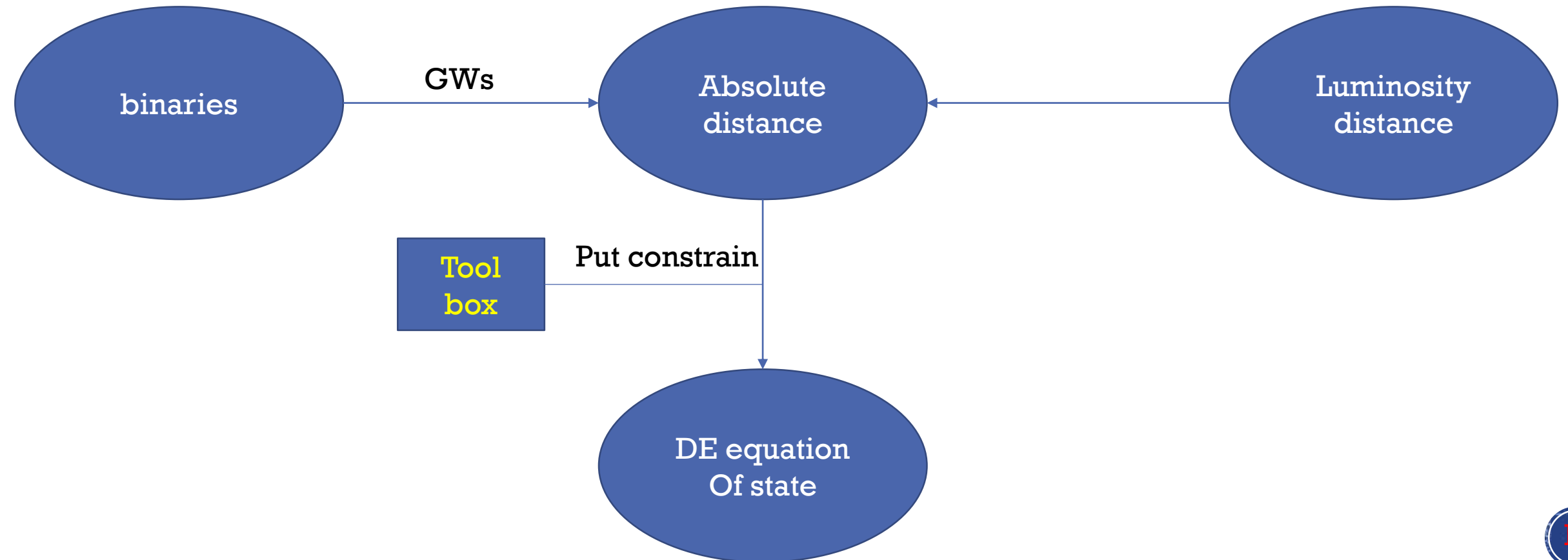
Identifying EM counter part

Getting redshift of producing source & dL

Gravitational lensing

Using mass distribution of binaries

# Getting redshift of producing source & dL



Luminosity distance

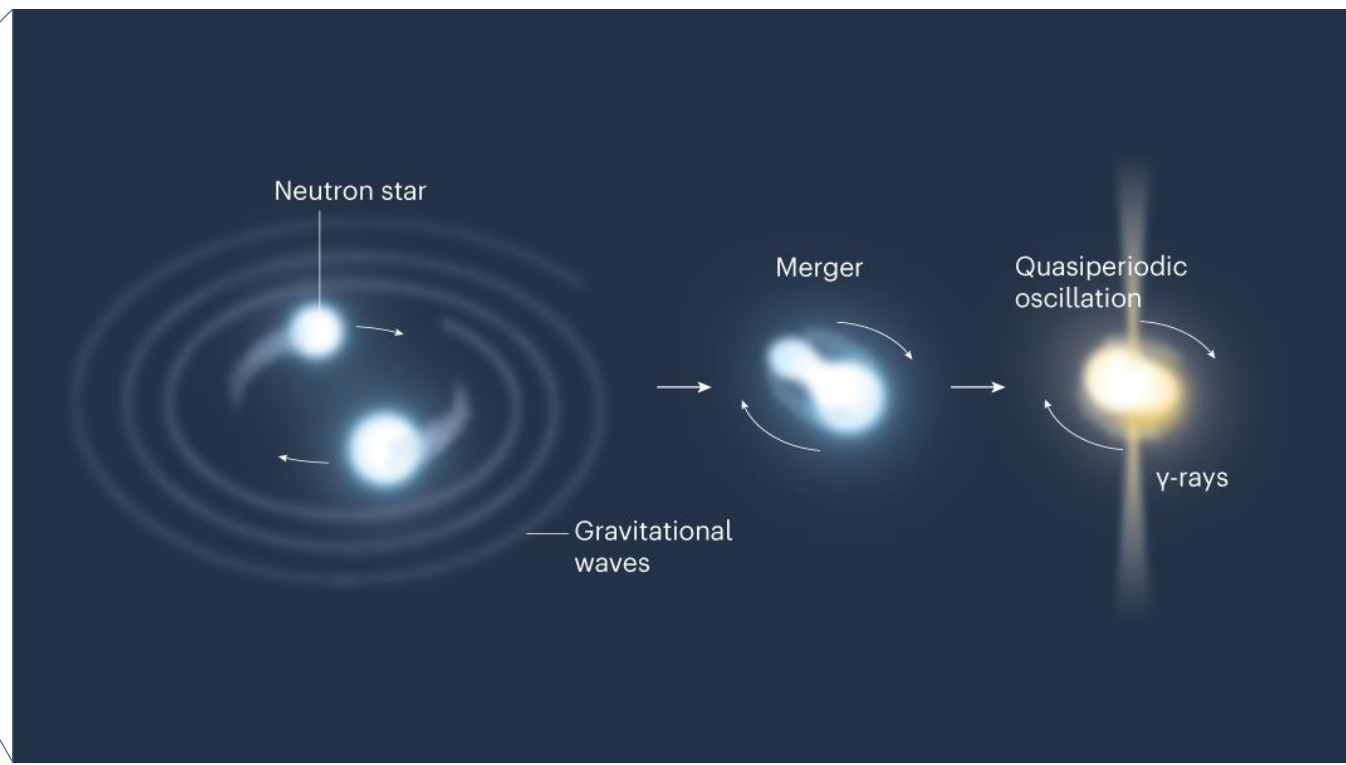


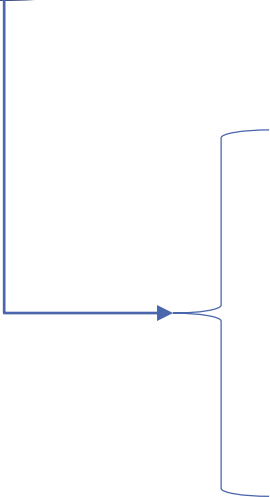
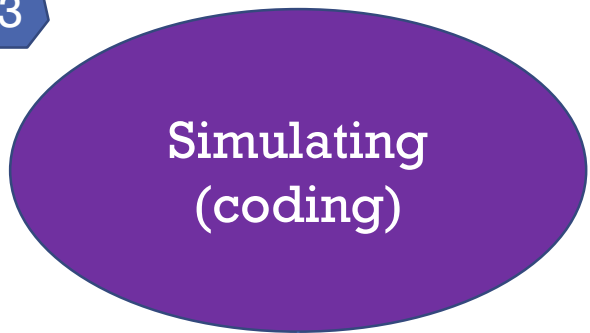
We need redshift !



- Galaxy catalog
- Neutron star mass distribution
- Tidal deformation of NS
- Binary merger of NS

Tool box





likelihood

posterior

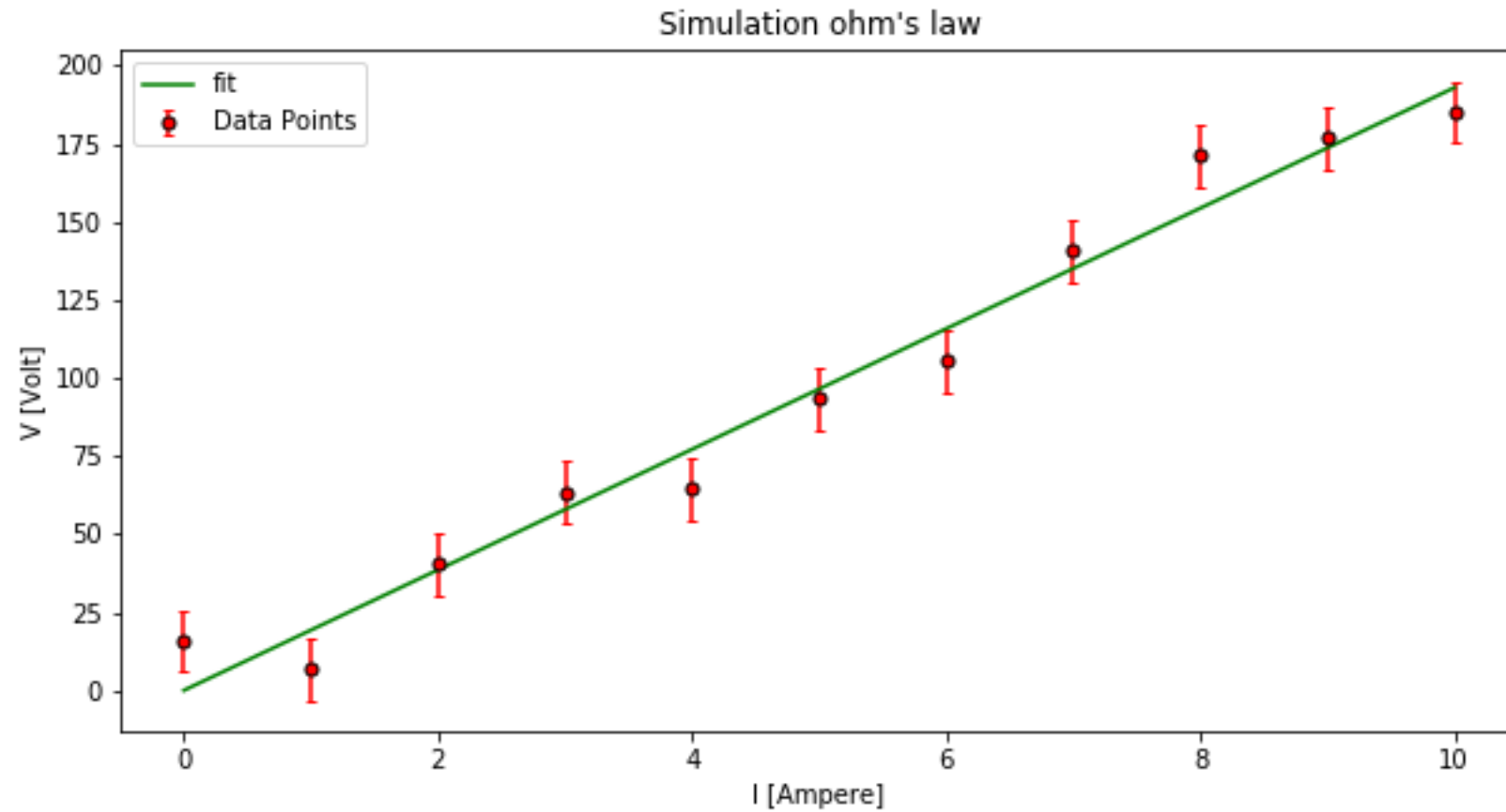
$$L(D|\Theta) \propto \exp(-\chi^2) \xrightarrow{\text{minimize}} \frac{\partial \chi^2}{\partial \Theta} = 0$$

$$\chi^2 = \frac{1}{N} (X_{th} - X_{obs})^2 \qquad V = R * I$$

$$\chi^2 = \frac{1}{N} (RI_I - V_i)^2$$

$$\frac{L(D|\Theta) p(\Theta)}{\int L(D|\Theta) p(\Theta)}$$

# What is it all about?!





سپاس از شما



## **INTENSIVE REPORT RELATED ARTICLES**

- [1]: <https://doi.org/10.1088/1475-7516/2023/03/047>
- [2]: <https://doi.org/10.3847/2041-8213/acda9a>
- [3]: <https://doi.org/10.1038/s42254-021-00303-8>
- [4]: <https://doi.org/10.3847/2041-8213/acdac6>
- [5]: arXiv:2306.16220v2
- [6] : <http://www.njp.org/doi:10.1088/1367-2630/7/1/204>
- [7] :10.1103/PhysRevResearch.4.013247
- [8]: EPJC-20-11-123
- [9]: <https://doi.org/10.1038/s42254-021-00303-8>
- [10] : <http://www.njp.org/doi:10.1088/1367-2630/7/1/204>
- [12]: EPJC-20-11-123
- [11] :10.1103/PhysRevResearch.4.013247
- [13] : arXiv:1608.08008v2
- [14] :arXiv:2202.09726v3
- [15] : arXiv:2009.09754v1
- [16] : arXiv:2109.07537v2